Cosmology of braneworlds *Basic principles & applications*

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TFA

Intuitevely: Our 4D spacetime is a timelike hypersurface in a 5D spacetime. The matter is confined to our spacetime

More precisely: $p\!\!-\!\!$ brane is a $p\!\!-\!\!$ dim spacelike surface in –dim spacetime , $D > p+1$

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Motivation

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- Closed string can move through the bulk, their spectrun contains gravitons

Recent history

Akrani–Hamed, Dimopoulose and Dvali braneworld – $d+d$ -dim flat bulk with d dim compactified (torus geometry)

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- Randall–Sundrum (RS) braneworld ^a flat braneworld embedded into 5D *Ad*S

Action

$$
S = -\int \mathrm{d}^5 x \sqrt{-g^{(5)}} \left(\frac{R}{2\kappa_5^2} + \Lambda_5 \right) + S_{\text{fields}} \;,
$$

 $\frac{1}{2} + \Lambda_5$ +
and cosn with κ_5 and Λ_5 bulk gravitational and cosmological constants

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Analogical to 4D case

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5D Einstein equation

$$
G_{ab} \equiv R_{ab} - \frac{1}{2}g_{ab}R = -\kappa_5^2 T_{ab} + g_{ab}\Lambda_5
$$

A static $SO(3)$.

1) symmetric metric

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ds^{2} = e^{2A(Z)} \eta_{\mu\nu} dx^{\mu} dx^{\nu} - dZ^{2}
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An expanding braneworld \mathbf{r}

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ds^{2} = a^{2}b^{2}(dt^{2} - dZ^{2}) - a^{2} \left(\frac{dr^{2}}{1 - Kr^{2}} + r^{2}d\theta^{2} + r^{2}\sin^{2}\theta d\phi^{2} \right)
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Embedding of FRW braneworld

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- Define \emph{jump} of a function f by

$$
[f] = \lim_{\epsilon \to +0} [f(Z + \epsilon) - f(Z - \epsilon)].
$$

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 $\frac{3}{1}$ Decompose stress–energy as $T_{ab} = \tau_{ab} - \sigma h_{ab}$

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- Decompose stress–energy as $T_{ab} = \tau_{ab} \sigma h_{ab}$
- $\frac{3}{4}$ Projected Einstein braneworld equations are

$$
G_{ab}^{(4)} = 8\pi G \tau_{ab} - \Lambda_4 h_{ab} + \kappa_5^4 \tau_{ab} - E_{ab}
$$

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Are given by

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where

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$$

\n
$$
\tau_{ab} = \frac{1}{12} \tau \tau_{ab} - \frac{1}{4} \tau_{ac} \tau_b^c + \frac{1}{8} h_{ab} \tau_{cd} \tau^{cd} - \frac{1}{24} \tau^2 h_{ab}
$$

\n
$$
E_{ab}
$$
 is electric Weyl tensor $E_{ab} = C_{acbd} n^c n^d$

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Higher energy corrections due to π_{ab} Main didifferences:

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The gravitational bulk influence due to The gravitational bulk influence due to E \mathbf{r}
	- Bianchi identity implies $\kappa^4_5\nabla^a\pi_{ab}=\nabla^aE$ \mathbf{r}

Embedding of Minkowski braneworld into 5D AdS

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- Metric is

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ds^{2} = e^{-2k|Z|} \eta_{\mu\nu} dx^{\mu} dx^{\nu} - dZ^{2} ,
$$

$$
\frac{\kappa_{5}^{2}}{6} \Lambda_{5}
$$

with
$$
k = \sqrt{-\frac{\kappa_5^2}{6} \Lambda_5}
$$

Coordinate singularities at $Z=\pm\infty$ \bullet

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- Intrinsic singularity on the branes
- Vanishing effective cosmological constants on the branes

Small fluctuations due to ^a point mass added on brane

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- The perturbed metric

$$
ds^{2} = [e^{-2k|Z|} \eta_{\mu\nu} + h_{\mu\nu}(x, Z)] dx^{\mu} dx^{\nu} - dZ^{2}
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Small fluctuations due to ^a point mass added on brane The perturbed metric

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 $\mathrm{d}s^2 = [\mathrm{e}^{-2\kappa|Z|}\eta_{\mu\nu} + h_{\mu\nu}(x,Z)]\mathrm{d}x^\mu\mathrm{d}x^\nu - \mathrm{d}Z$
Equation for $h_{\mu\nu}$ admits separation of variables, $\mu\nu(x^\rho,Z) = \psi(Z) \Phi(x^\rho)$ with
 $\partial^\mu \partial_\mu \Phi(x^\rho)$

$$
\partial^{\mu}\partial_{\mu}\Phi(x^{\rho}) = -m^2\Phi(x^{\rho})
$$

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In cylindrically symmetric case in zero–mass mode

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\Psi(r) = -\frac{B}{r}
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 $\frac{r}{2}$ ji Kaluza–Klein non–zero modes give r^{-3} corrections
 ϵ

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	- Electroweak spontaneous symmetry breaking at energy scales $M_{EW}\sim 10^3~{\rm GeV}$

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	- Electroweak spontaneous symmetry breaking at energy scales $M_{EW}\sim 10^3~{\rm GeV}$
	- Stringy effects at scales $M_{\rm Pl}$ $\sim 10^{19}$ GeV
	- A huge gap between M_{Planck} and M_{EW} , 16 orders (!)

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	- Relation between a physical mass m (observed on the 3-brane) and its corresponding bulk mass $m_{\rm 0}$

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	- Relation between a physical mass m (observed on the 3-brane) and its corresponding bulk mass $m_{\rm 0}$

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it is sufficient $kr_c \approx 50$ in order to obtain the hierarchy

FRW braneworld

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$$
{\rm d}s^2 = a^2b^2({\rm d}t^2-{\rm d}Z^2) - a^2\left(\frac{{\rm d}r^2}{1-Kr^2} + r^2{\rm d}\theta^2 + r^2\sin^2\theta{\rm d}\phi^2\right)
$$

Basic principles April 15, 2004 -
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$$
H^{2} = \frac{8\pi G}{3} \left[\rho_{M} + \frac{\rho_{M}^{2}}{2\sigma} \right] + \frac{\Lambda_{4}}{3} + \frac{\mu}{a^{4}}
$$

$$
\frac{dH}{d\tau} = -4\pi G(\rho_{M} + p_{M}) \left[1 + \frac{\rho_{M}}{\sigma} \right]
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- \overline{M} is mass depaity
- Braneworld versions of Friedmann and Raychaudhuri $equating$ ns
- **Dark radiation** term

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H^{2} = \frac{8\pi G}{3} \left[\rho_{M} + \frac{\rho_{M}^{2}}{\rho_{M}} \right] + \frac{\Lambda_{4}}{3} + \frac{\mu}{a^{4}}
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- \overline{M} is mass density
- Braneworld versions of Friedmann and Raychaudhuri equations
- **Dark rad***i***ation term**
- Quadratic term ρ_M

$$
H^{2} = \frac{8\pi G}{6} \left[\rho_{M} \left(\frac{\rho_{M}^{2}}{2\sigma} \right) + \frac{\Lambda_{4}}{3} + \frac{\mu}{a^{4}} \right]
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- \overline{M} is mass density
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- **Dark radiation** term
- For $\rho_M\mathord{>}\mathord{>}\sigma$ it is $H\sim\rho$

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orld *versions of Friedman and Ray*

- \overline{M} is mass density
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- **Dark radiation term**
- For $\rho_M\!\!>\!\not\! k\sigma$ it is $H\sim$

For $\rho_M<<\sigma$ it is $H\sim\sqrt{\rho_M}$ – classical scenario

According to classical cosmology, the FRW universe expands into "nothing"

According to braneworld scenario, the FRW brane expands into 5D AdS spacetime